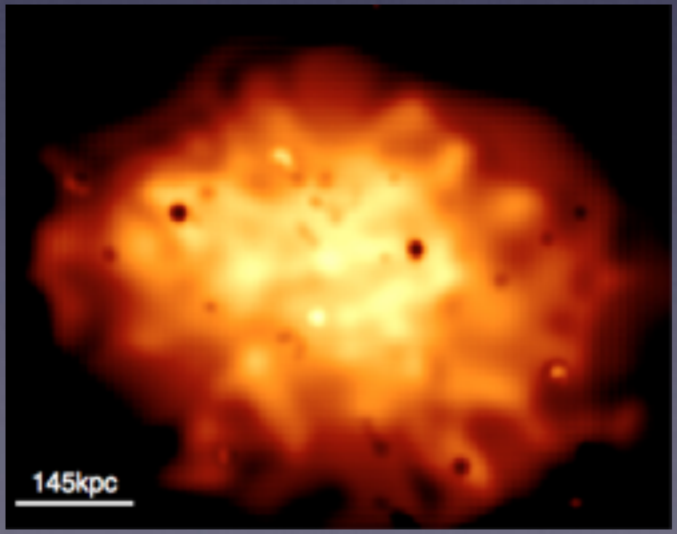
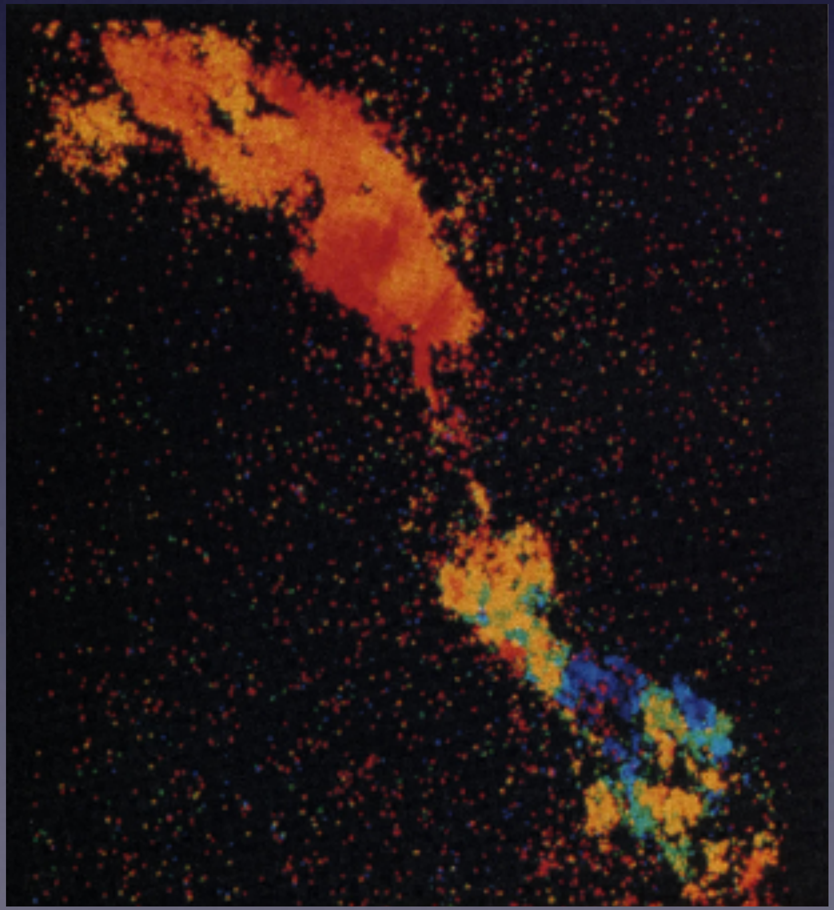
# Collisionless plasma dynamo

François Rincon (IRAP Toulouse)

with Francesco Califano (U. Pisa),
Alex Schekochihin (Oxford), F. Valentini (U. Calabria)

Acknowledgements:
S. Cowley, M. Kunz, C. Cavazzoni

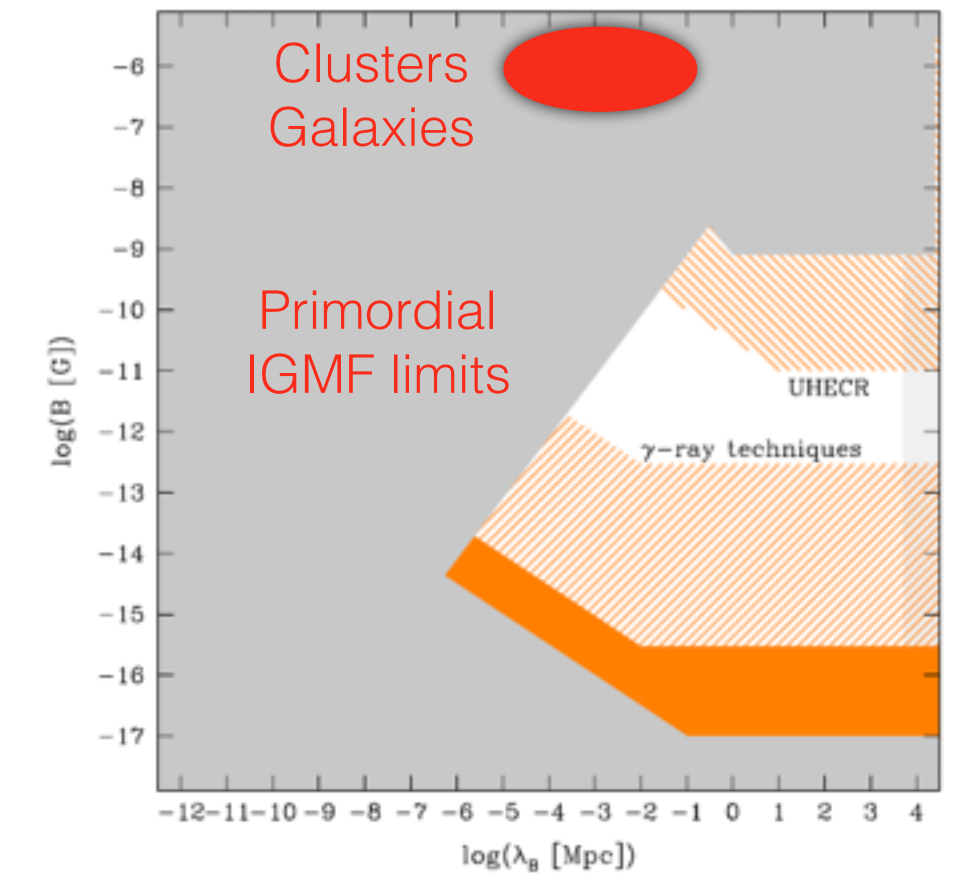


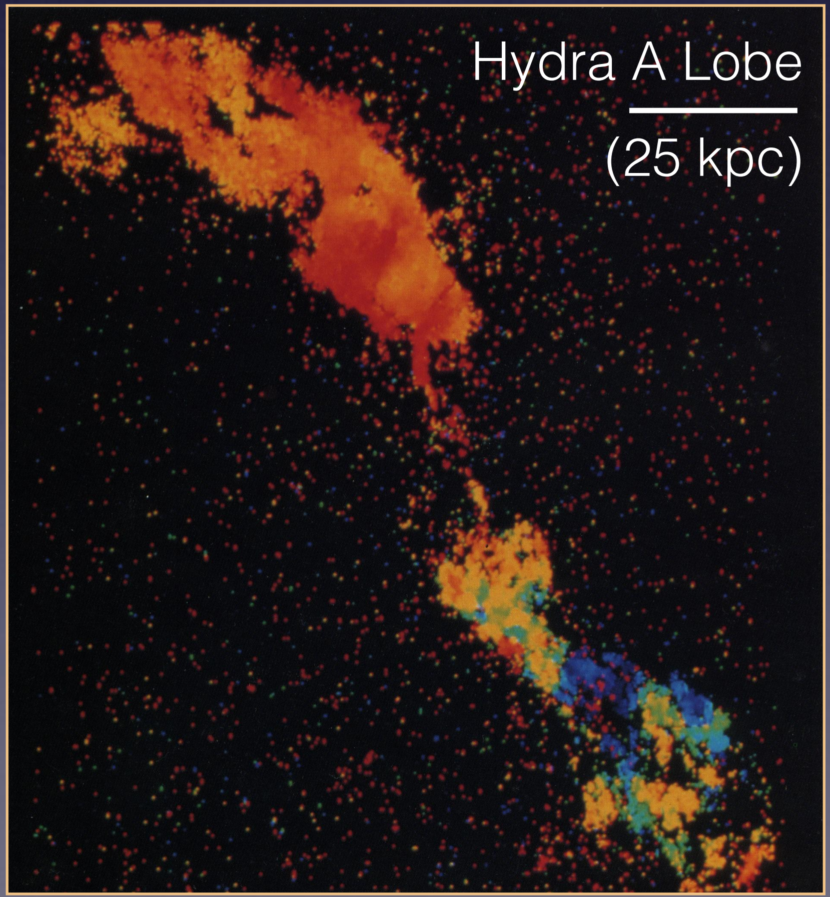
Vlasovia, June 2016

# Cosmic magnetogenesis

- How are magnetic fields generated on cosmic scales ?
  - Magnetic seeds in the early Universe: $10^{-21}(-10^{-9}\ ?)$ G
  - ICM fields: 1-40 μG at fairly large (~ 1-10 kpc) scales
  - Constraint: 5-15 fold increase on a few Gyr




Durrer & Neronov, A&A Rev. 2013



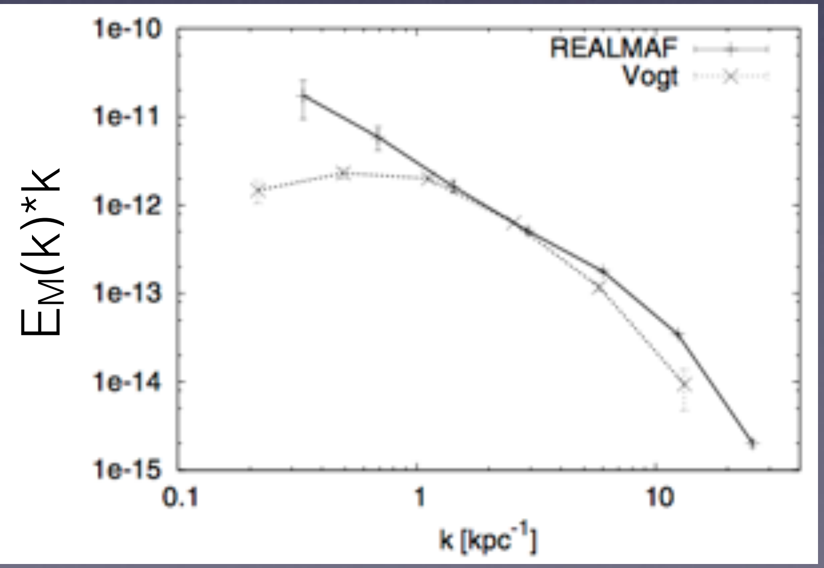

Taylor & Perley, ApJ 1993

# ICM magnetic fields

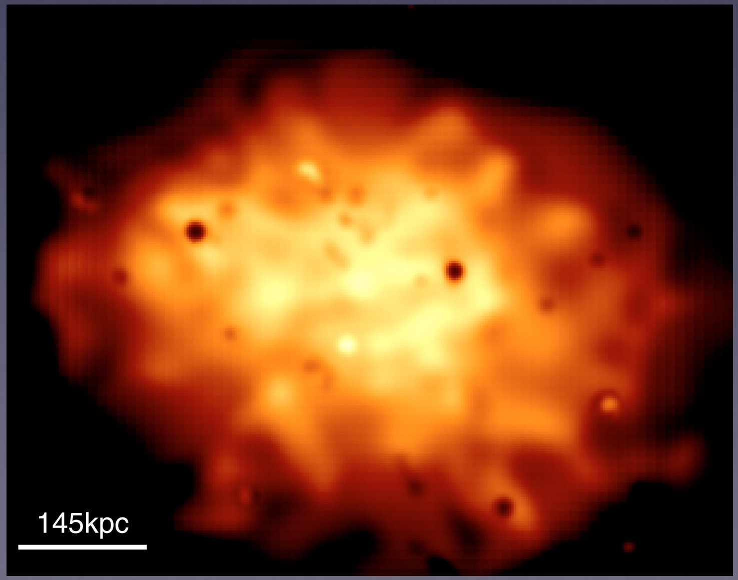
- How do you make microGauss fields at 1-100 kpc scales ?

- Different processes invoked
  - Magnetization via galactic outflows and jets
  - Collisionless shocks in ICM / filaments
  - Dynamo effect throughout cosmic times

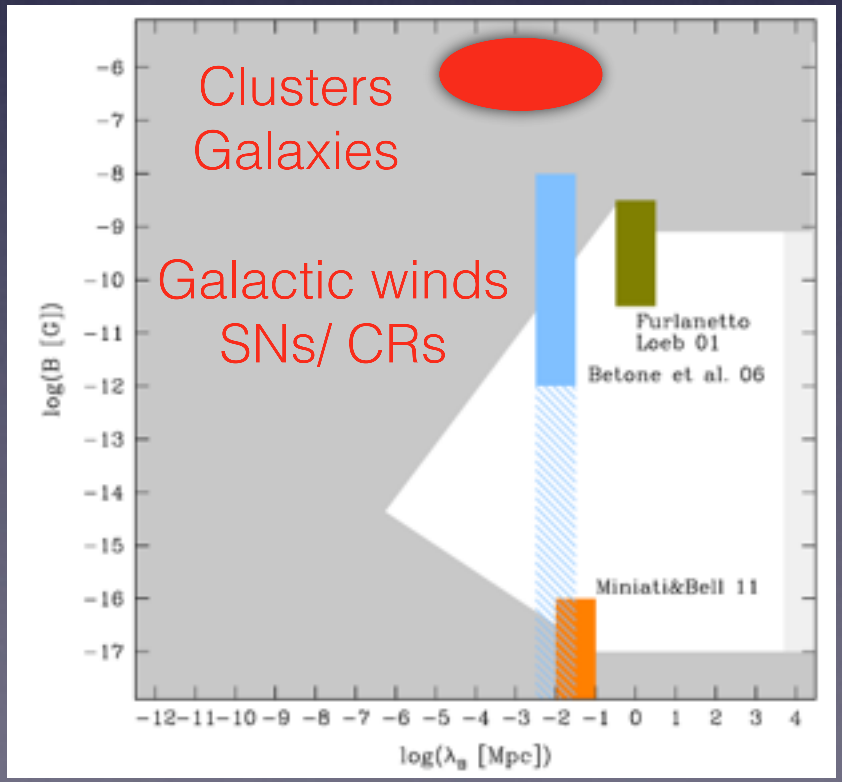
- Is turbulence (T~10-100 Myr) in the ICM or filaments a good dynamo ?




Kuchar & Ensslin, A&A 2011



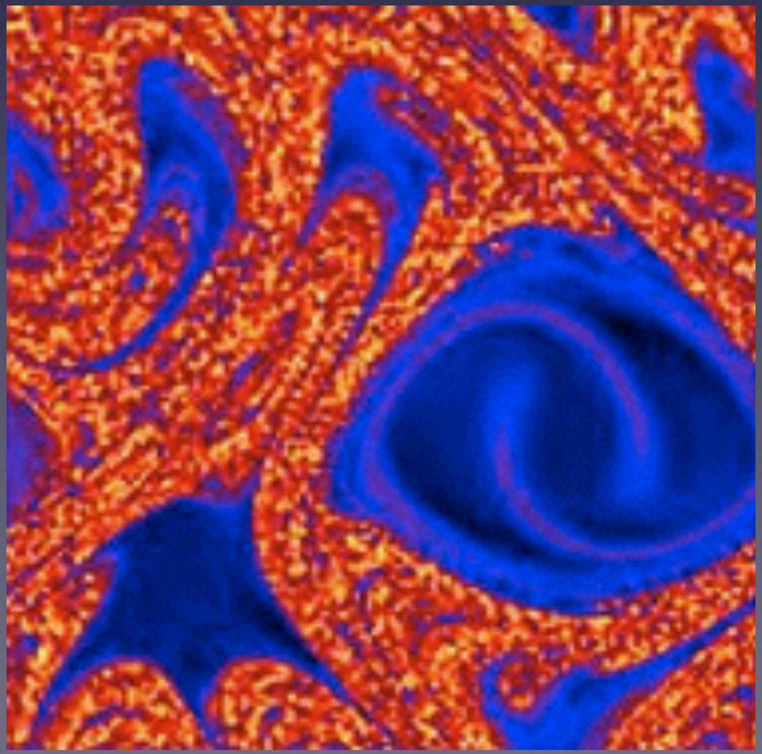

Schueker et al., A&A 2004

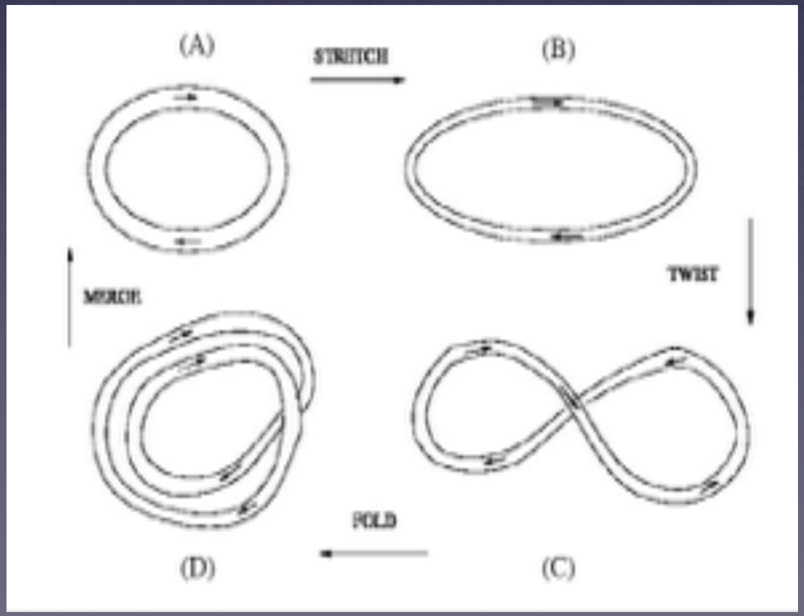


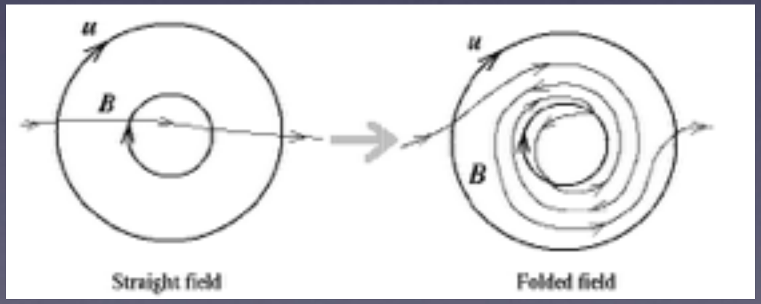

Durrer & Neronov, A&A Rev. 2013

# Turbulent “small-scale” dynamo

- Homogeneous, isotropic, non-helical, incompressible, chaotic flow of conducting fluid is a dynamo flow
  - Batchelor-Moffatt-Zeldovich’s stretch-fold mechanism
  - All you need is a smooth 3D chaotic flow, viscous flow can do the job





# First evidence in 3D MHD simulations

**Helical and Nonhelical Turbulent Dynamos**

M. Meneguzzi
*Centre National de la Recherche Scientifique and Section d'Astrophysique, Division de la Physique, Centre d'Etudes Nucléaires de Saclay, F-91191 Gif-Sur-Yvette, France*

and

U. Frisch
*Centre National de la Recherche Scientifique, Observatoire de Nice, F-06007 Nice, France*

and

A. Pouquet[a]
*Centre National de la Recherche Scientifique, Observatoire de Meudon, F-92190 Meudon, France*

(Received 13 April 1981)

Direct numerical simulations of three-dimensional magnetohydrodynamic turbulence with kinetic and magnetic Reynolds numbers up to 100 are presented. Spatially intermittent magnetic fields are observed in a flow with nonhelical driving. Small-scale helical driving produces strong large-scale nearly force-free magnetic fields.

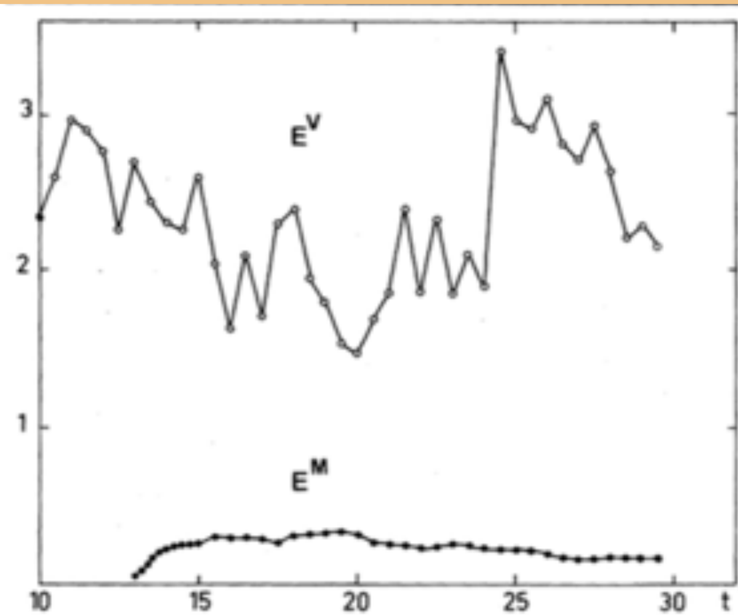



FIG. 1. Turbulent dynamo with nonhelical driving. Temporal variation of kinetic ($E^V$) and magnetic ($E^M$) energy. Reynolds numbers are $R^V = R^M \approx 100$. The time unit is the eddy-turnover time $l_0/v_0$.

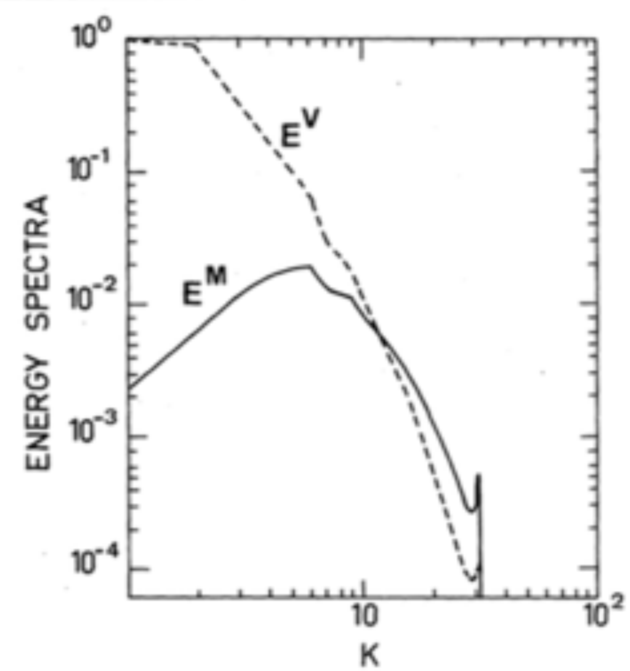
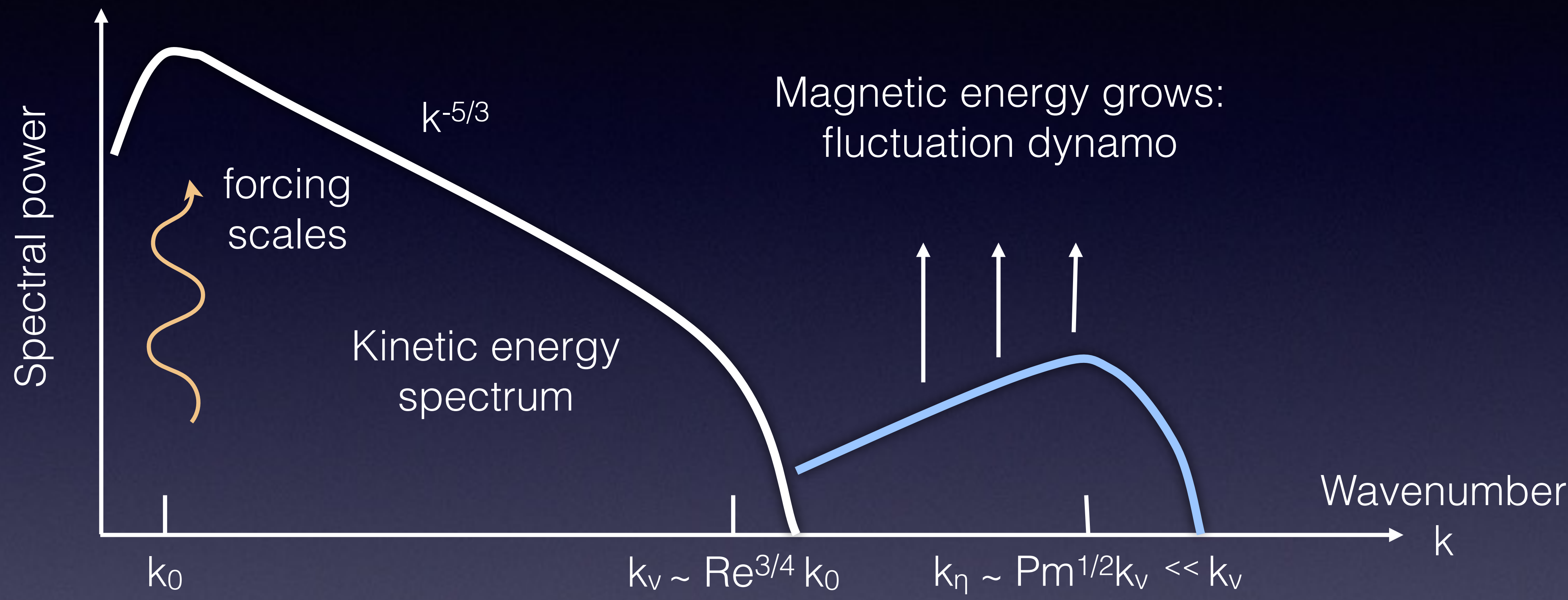



FIG. 2. Kinetic ($E^V$) and magnetic ($E^M$) energy spectra at $t = 27$. Nonhelical dynamo with $R^V = R^M \approx 100$.

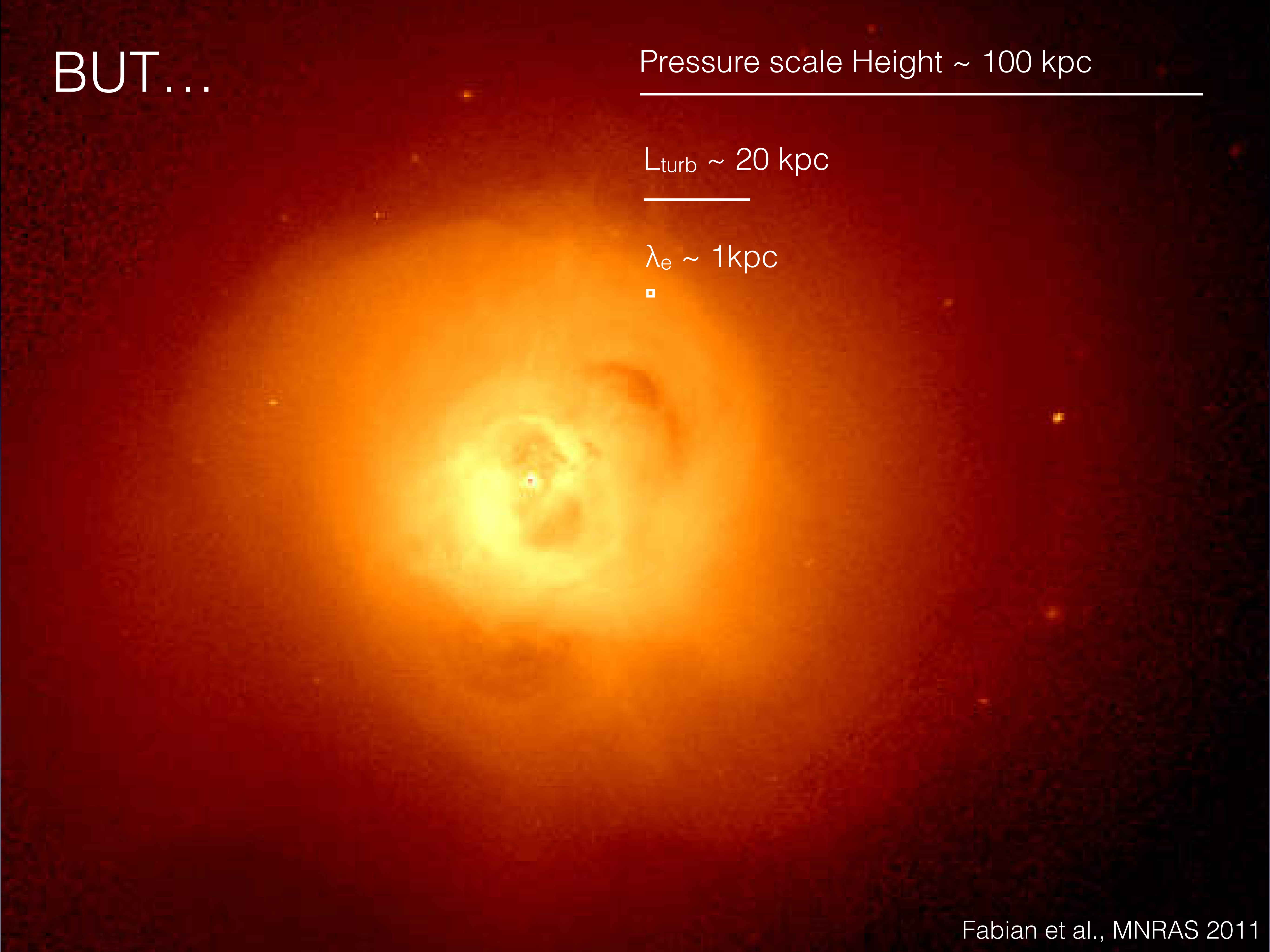
# Large magnetic Prandtl number regime

- In such a fluid, the dynamo field grows at small scales




- Naive ICM "MHD" parameters
  - Collisional viscosity estimate: $Re \sim UL/\nu \sim$ 10-100
  - Spitzer conductivity: $Rm \sim UL/\eta \sim 10^{29}$ or more
  - Magnetic Prandtl number $Pm \sim \nu/\eta \sim 10^{28\text{-}30}$

# BUT…



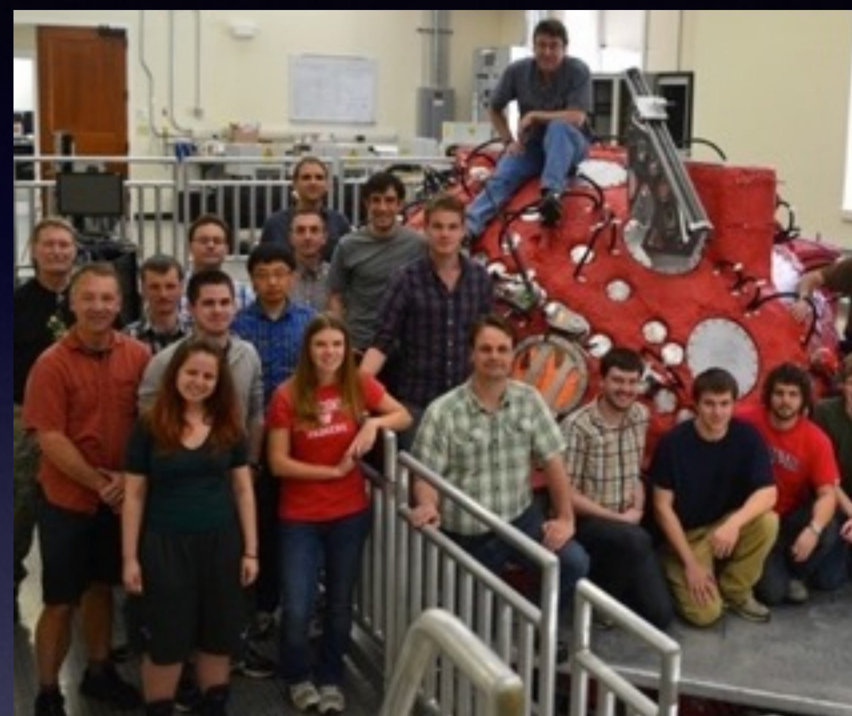

Fabian et al., MNRAS 2011

# What about weakly-collisional plasmas ?

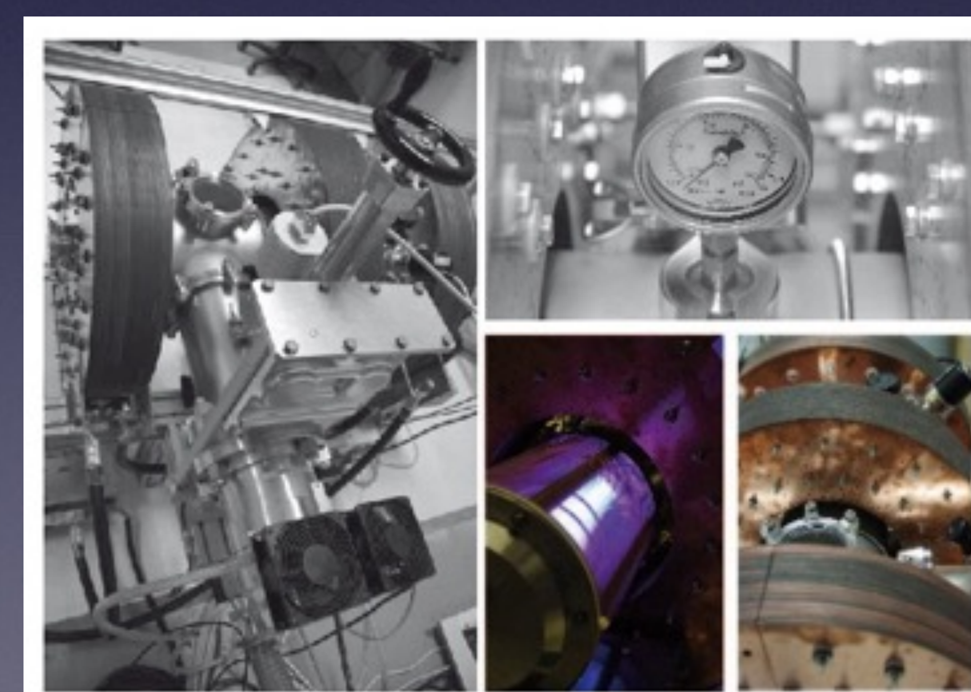
- So far, dynamo has only been demonstrated in MHD fluids
  - Many high-energy astrophysical plasmas are not MHD fluids
- ICM plasma regime
  - Dynamical/injection scales ~ $10^{17-18}$ km ~ 10 - 100 kpc (T~10-100 Myr)
  - Mean free path ~$10^{16-17}$ km ~ 1-10 kpc
  - Larmor radii ~ $10^4$ km
- Coupled "fluid-" and "kinetic-scale" phenomena
  - Large-scale dynamics: MTI, HBI, AGN, mergers, dynamo ?
  - Collisionless damping, magnetization effects (pressure anisotropies)

# Plasma dynamo: an experimental quest in progress

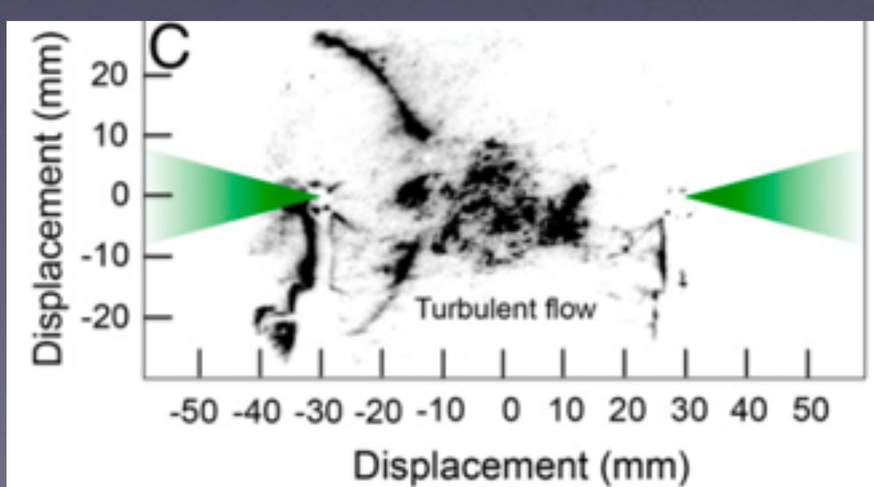
Madison Plasma Dynamo Experiment @U. Wisconsin



Turbulent Plasma experiment
@ ENS Lyon

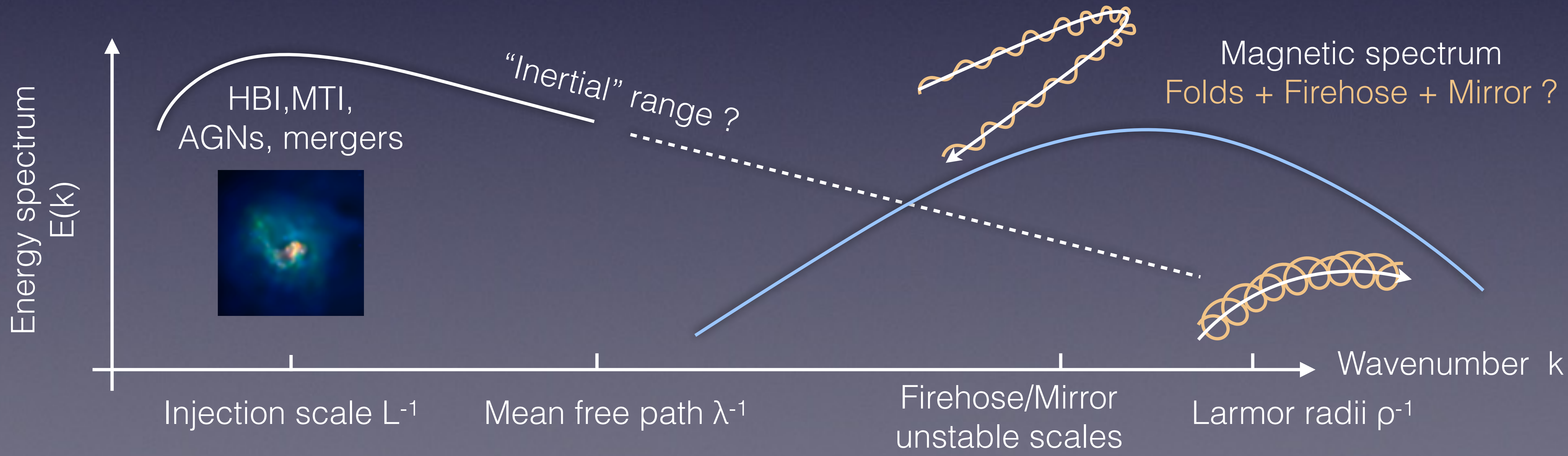



Oxford Laser Plasma group
(Gregori, Meinecke et al., PNAS 2015)




# Collisionless plasma dynamo problem

- The most efficient eddies are the smallest, fastest ones
  - In the ICM, such plasma motions are weakly collisional
- Plasma is magnetised well below equipartition (ICM: $10^{-13}$ G)
  - Field-stretching motions (= dynamo !) generate pressure anisotropy
  - Pressure-anisotropy driven instabilities !

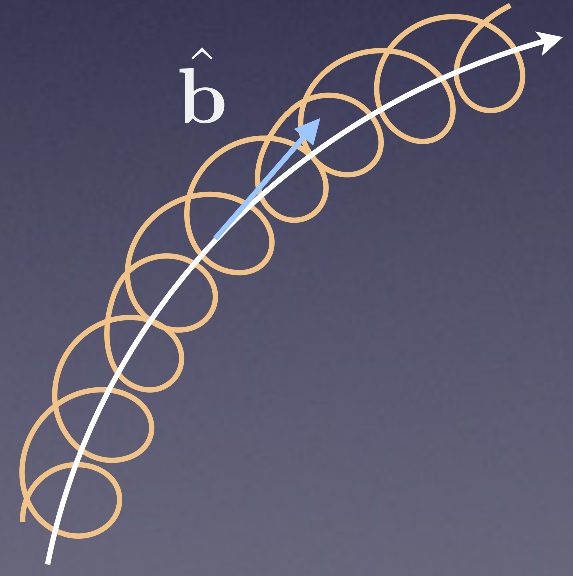


# Pressure anisotropy generation

- In a magnetized, weakly collisional plasma
  - The pressure is an anisotropic tensor with respect to the direction of B
  - $\mu_s = m_s v_\perp^2 / 2B$ is almost conserved
- Large-scale, field-stretching motions generate pressure anisotropy
  - Collisions tend to relax it

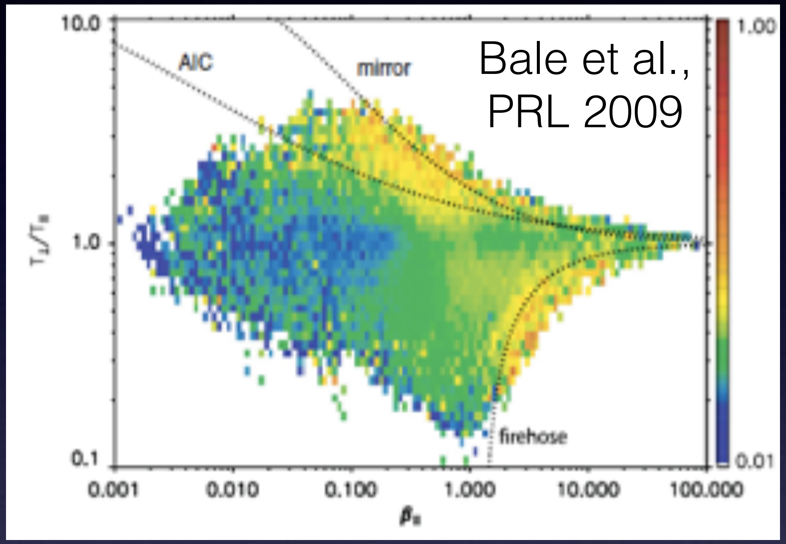
$$\frac{1}{p_\perp}\frac{\mathrm{d}p_\perp}{\mathrm{d}t} \sim \frac{1}{B}\frac{\mathrm{d}B}{\mathrm{d}t} - \nu_{ii}\frac{p_\perp - p_\parallel}{p}$$

$$\frac{1}{B}\frac{\mathrm{d}B}{\mathrm{d}t} = \hat{\mathbf{b}}\hat{\mathbf{b}} : \nabla\mathbf{u}$$



# Pressure anisotropy-driven instabilities

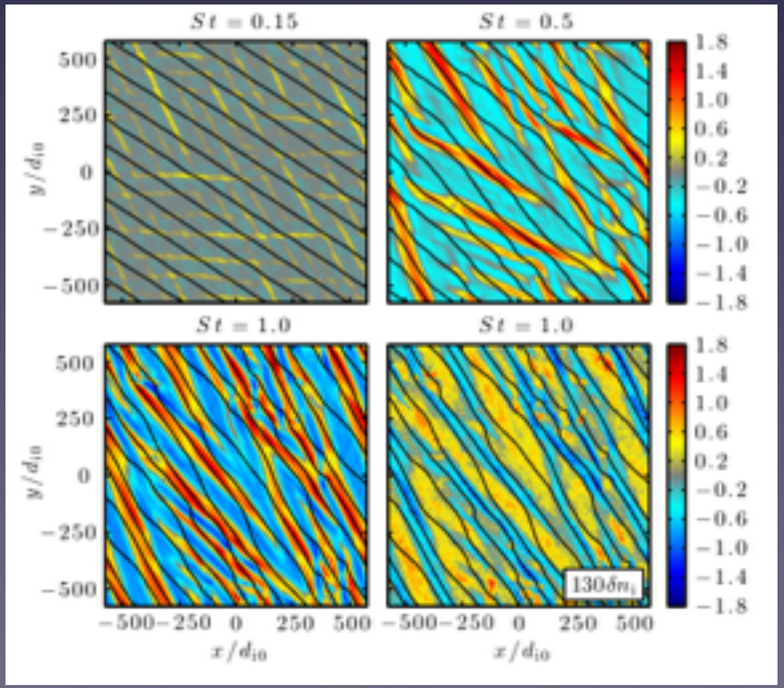
- $\mu = mv_\perp^2/2B$ conservation implies kinetic instability everywhere
  - local increase of |B| —> increase of $p_\perp$
    - mirror instable $\dfrac{p_\perp - p_\|}{p_\perp} > 1/\beta$
  - local decrease of |B| —> decrease of $p_\perp$
    - firehose instable $\dfrac{p_\perp - p_\|}{p_\perp} < -2/\beta$





- Small, fast scales
  - ICM: $\rho_i$ ~ $10^4$ km, $\Omega_i^{-1}$ ~ second
- Feedback non-linearly on “fluid” scales

Scheckochihin et al, ApJ 2005, Schekochihin et al., PRL 2008; Rosin et al., MNRAS 2011; Rincon et al., MNRAS 2015



Kunz et al., PRL 2014

# Collisionless plasma dynamo problem(s)

- Unmagnetized problem: $\rho_i/L > 1$
  - Is a collisionless, unmagnetized 3D chaotic flow of plasma a good dynamo ?
- Magnetized problem: $\rho_i/L < 1$
  - How do pressure-anisotropy kinetic instabilities interfere with magnetic growth ?
- Annoying “details”
  - Dynamo is a fundamentally 3D process in physical space (Cowling)
  - No rigid “guide” field here: kinetic description “3V” in velocity space
- Modelling requires 3D-3V simulations (+time integration !)
  - Very costly: O($10^6$-$10^7$ CPU hours) per simulation
  - Use simplest possible appropriate kinetic model

# Forced hybrid Vlasov-Maxwell system

- Kinetic, collisionless ions (initially Maxwellian)

$$\frac{\partial f_i}{\partial t} + \mathbf{v} \cdot \nabla f_i + \left[ \frac{e}{m_i} \left( \mathbf{E} + \frac{\mathbf{v} \times \mathbf{B}}{c} \right) + \frac{\mathbf{F}}{m_i} \right] \cdot \frac{\partial f_i}{\partial \mathbf{v}} = 0$$

- Isothermal, fluid massless electrons

$$\mathbf{E} = -\frac{T_e \nabla n_e}{e n_e} - \frac{\mathbf{u}_e \times \mathbf{B}}{c} + \frac{4\pi\eta}{c^2}\mathbf{j}$$

$$\mathbf{u}_e = \mathbf{u}_i - \mathbf{j}/(e n_e) \qquad \mathbf{j} = (c/4\pi)\,\nabla \times \mathbf{B}$$

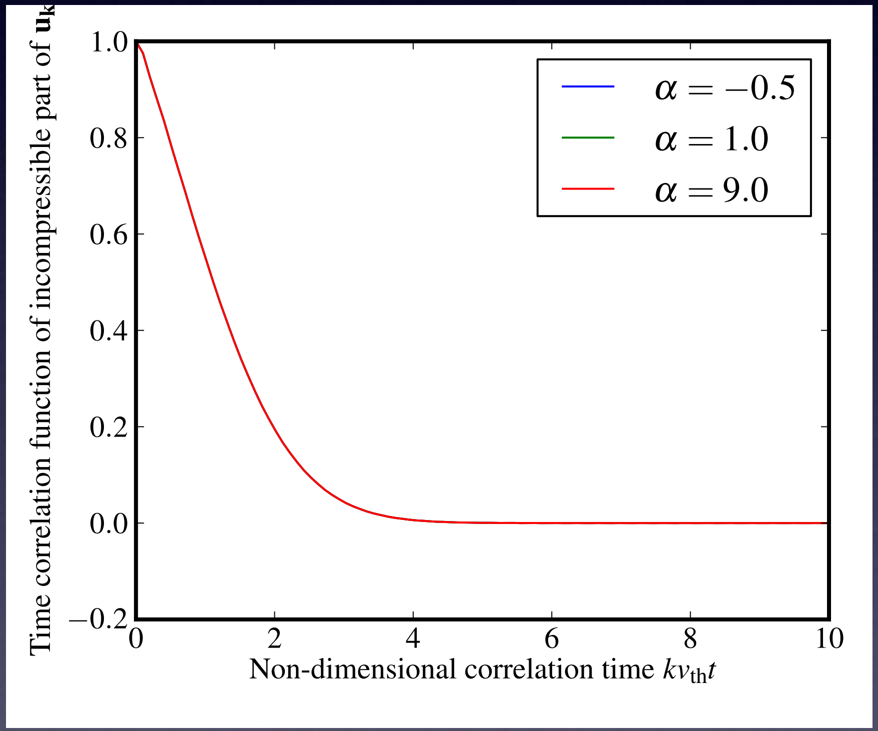
- Quasi-neutrality: $n_e = n_i$

$$\nabla \cdot \mathbf{B} = 0$$

- Maxwell-Faraday: $\displaystyle \frac{\partial \mathbf{B}}{\partial t} = -c\,\nabla \times \mathbf{E}$

# Collisionless flow forcing

- $\delta$-correlated-in-time large-scale forcing in kinetic ion equation
- In the unmagnetized regime, flow statistics controlled by phase-mixing (collisionless damping)
  - Flow correlation time is $(k_f v_{\mathrm{th}i})^{-1}$, a factor Mach number smaller than the turnover time
  - the flow is effectively highly viscous




$$\langle F_{\mathbf{k},i}(t) F^*_{\mathbf{k},j}(t')\rangle = \chi(k)\,\delta(t-t')\,(\delta_{ij} - k_i k_j/k^2)$$

$$\langle u_{\mathbf{k},i}(t) u^*_{\mathbf{k},j}(t')\rangle = \frac{\chi(k)}{8\pi k^2}\left(\delta_{ij} - \frac{k_i k_j}{k^2}\right)\int_{-\infty}^{\infty} d\omega\, e^{-i\omega(t-t')}\left| Z\left(\frac{\omega}{k v_{\mathrm{th}i}}\right)\right|^2$$

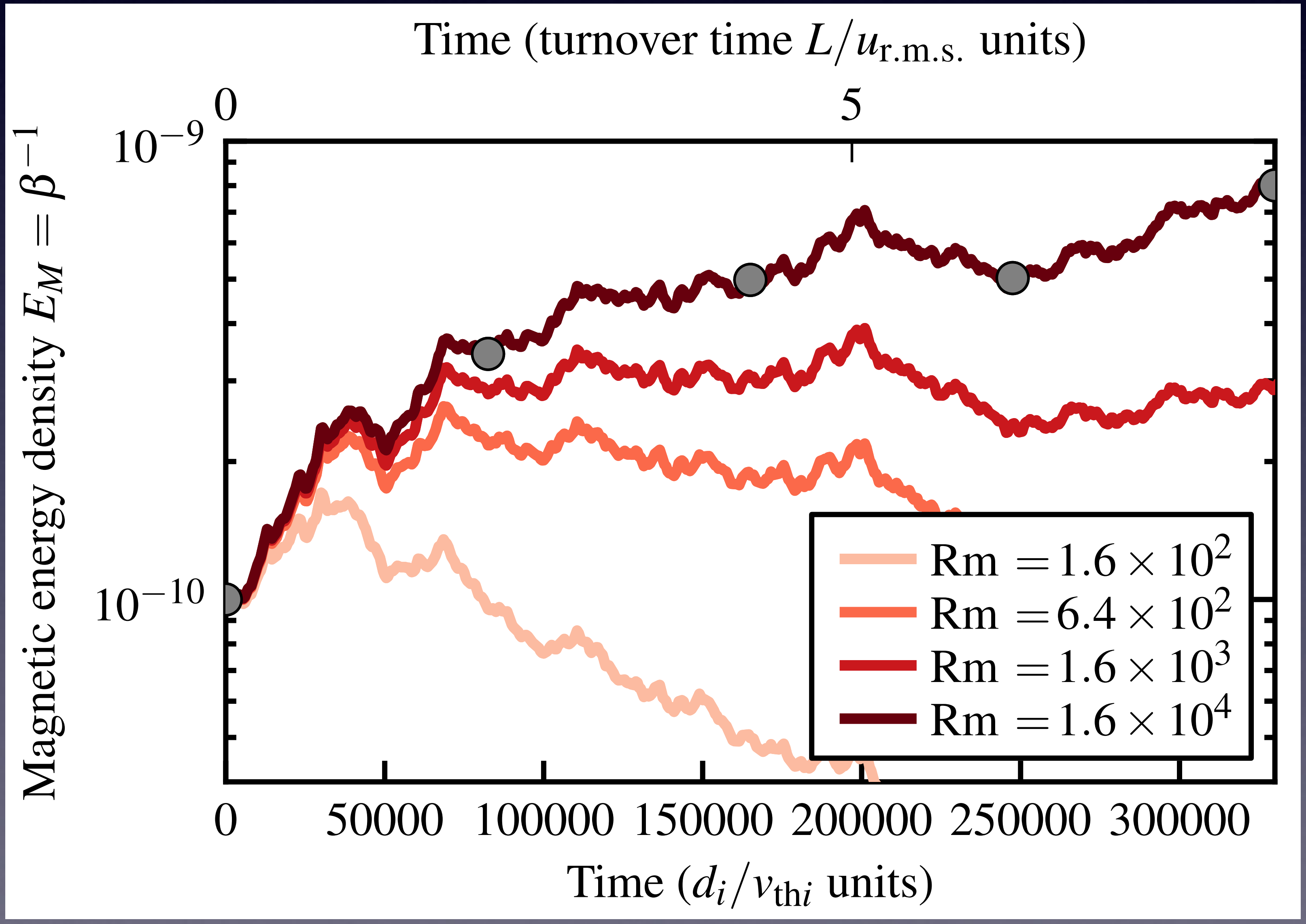
- Smooth, large-scale, chaotic, subsonic, finite-amplitude flow

# Dynamo simulations setup

- Solve hybrid Vlasov-Maxwell in 3D-3V with Eulerian code
  - 3D periodic, phase-space dimensions: $L = 2000\pi d_i\,,\;\; v_{\rm max} = \pm 5 v_{{\rm th}i}$
  - Resolution: $64^3$ (physical space) x $51^3$ (velocity space) (Valentini et al., JCP 2007)
- Incompressible, isotropic, non-helical delta-correlated forcing
  - $k_f = 2\pi/L\,,$ injected power $\varepsilon = 3\times 10^{-5}\, n_{i0} m_i v_{{\rm th}i}^3/d_i$
  - Box-scale, collisionless chaotic flow $u_{\rm r.m.s.} \sim 0.2\, v_{{\rm th}i}$
- Initial conditions
  - Isotropic ion Maxwellian, $T_e$=$T_i$
  - Magnetic seed in wavenumber range $[2\pi/L, 4\pi/L]$
    - No guide/mean field !
    - Magnetic energy measured as inverse of plasma $\beta = 8\pi\, n_{\rm i0} T_i/B^2_{\rm r.m.s.}$

# Unmagnetized regime

- Four simulations with same initial field and flow history, but different magnetic diffusivity $\eta$

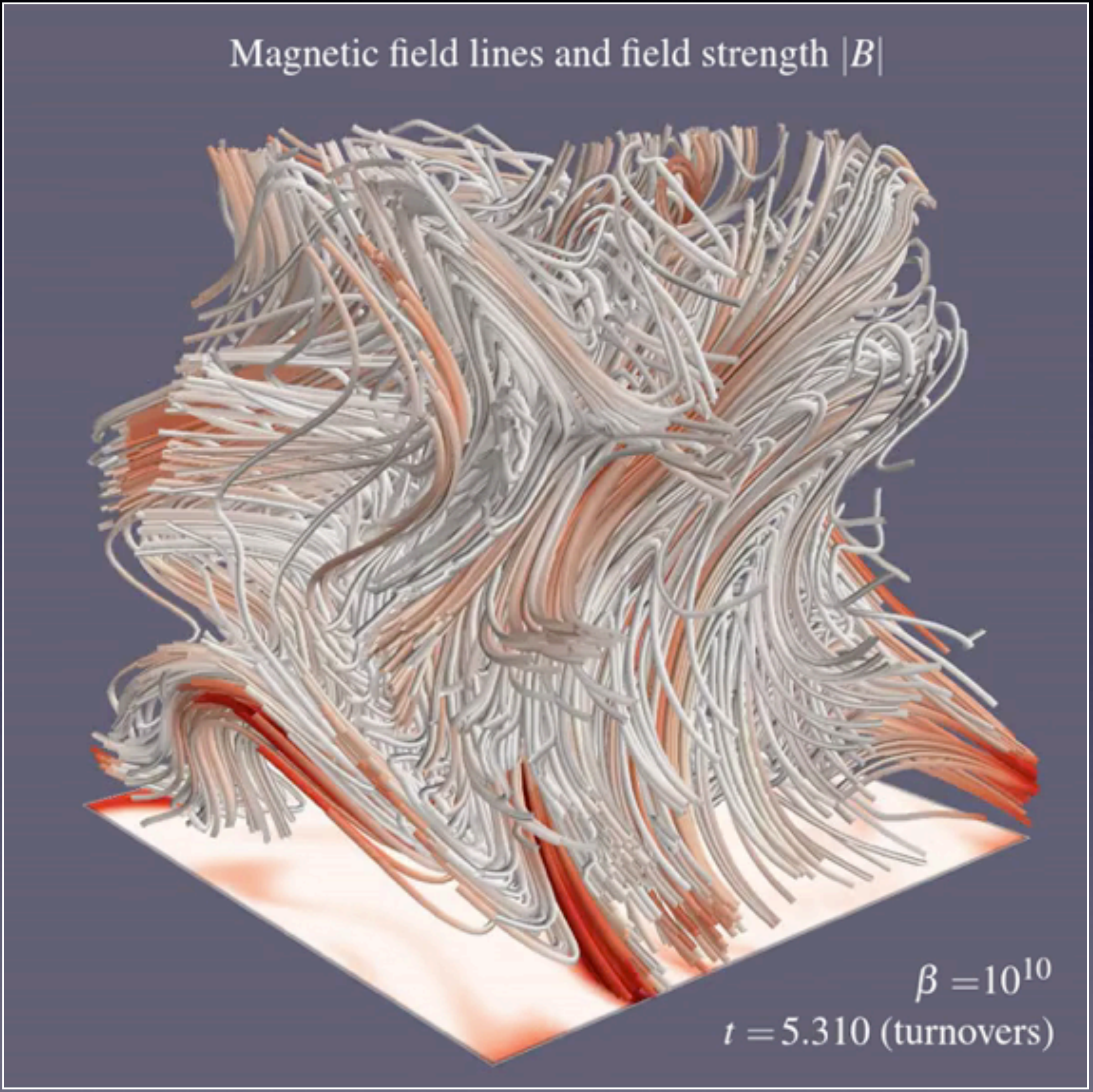



$$\beta = 10^{10}$$

$$\rho_i / L = 16$$

$$\mathrm{Rm} = \frac{u_{\mathrm{r.m.s.}}}{\eta k_f}$$

# Unmagnetized regime: growing case

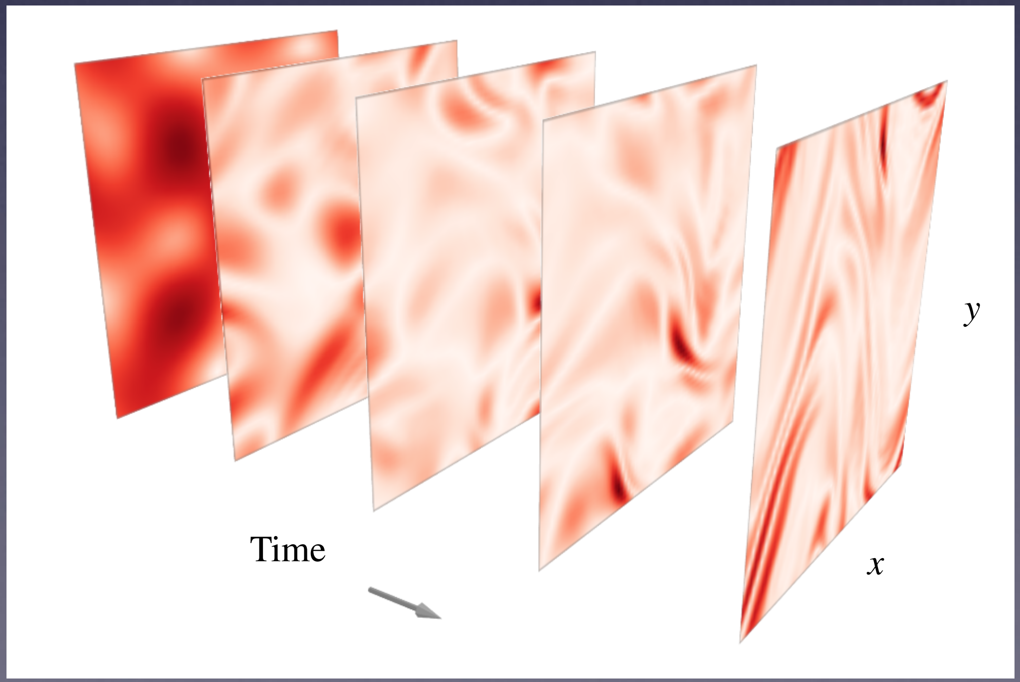


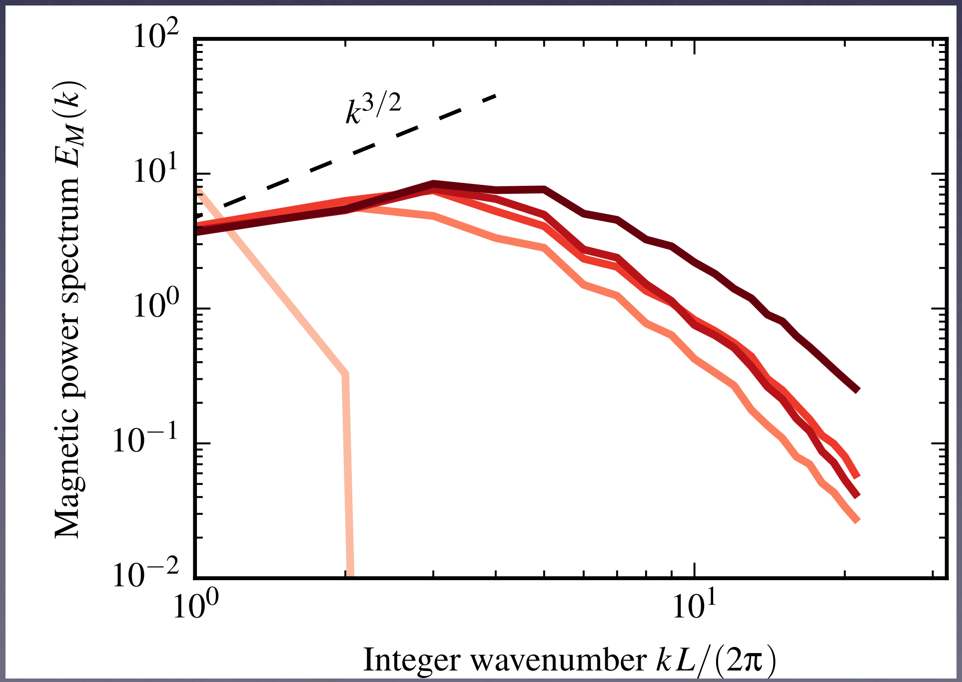

$\beta = 10^{10}$

$\rho_i / L \simeq 16$

# Small-scale dynamo

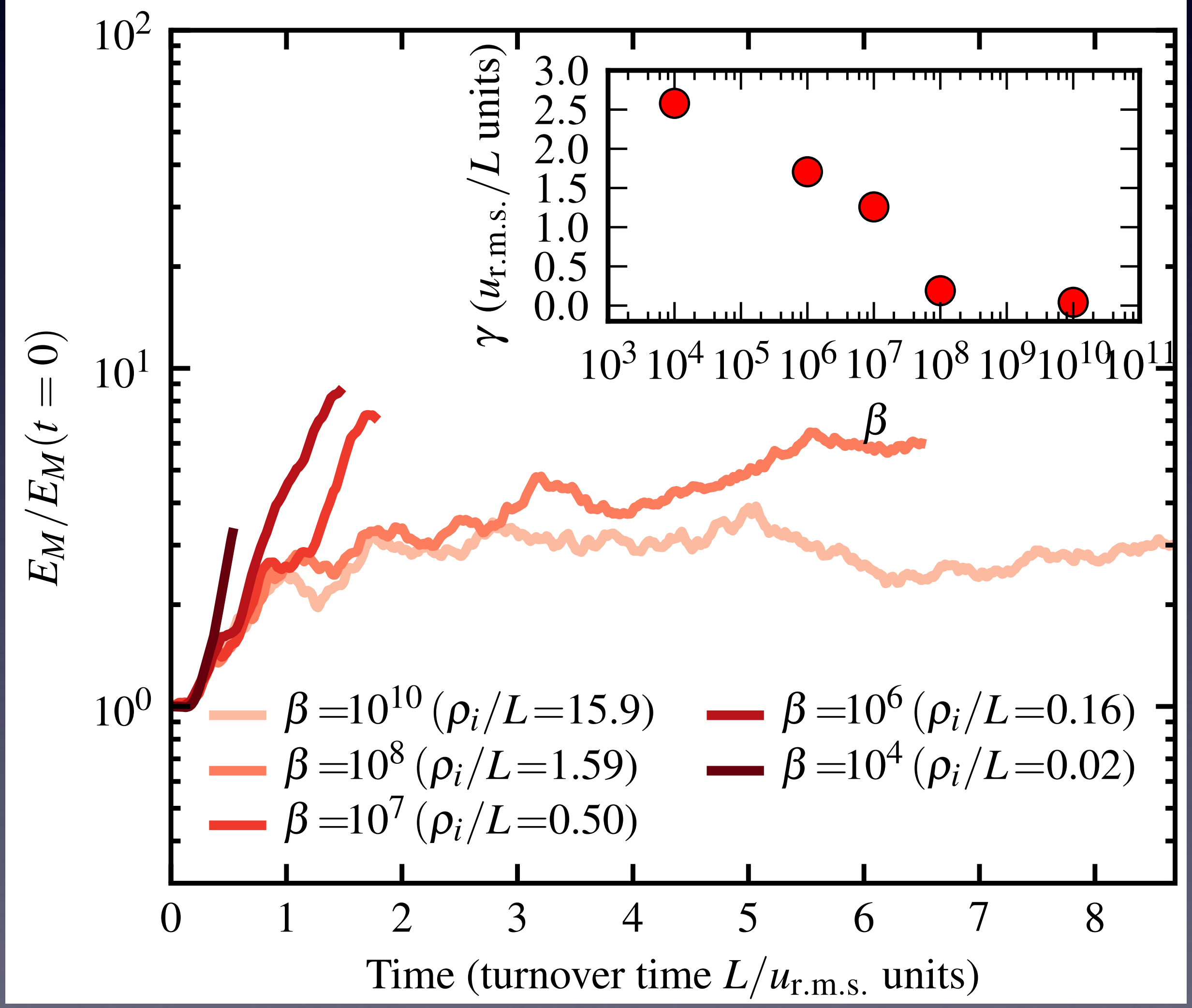
- Dynamo relies on chaotic stretching and folding of field lines
  - Folded field structure
  - Spectral evolution consistent with the formation of a Kasantsev spectrum
- Critical Rm larger than in MHD
  - Interpreted as a small flow correlation time effect
  - Energy growth rate ~ 0.15 turnover rate for Rm ~ 15000



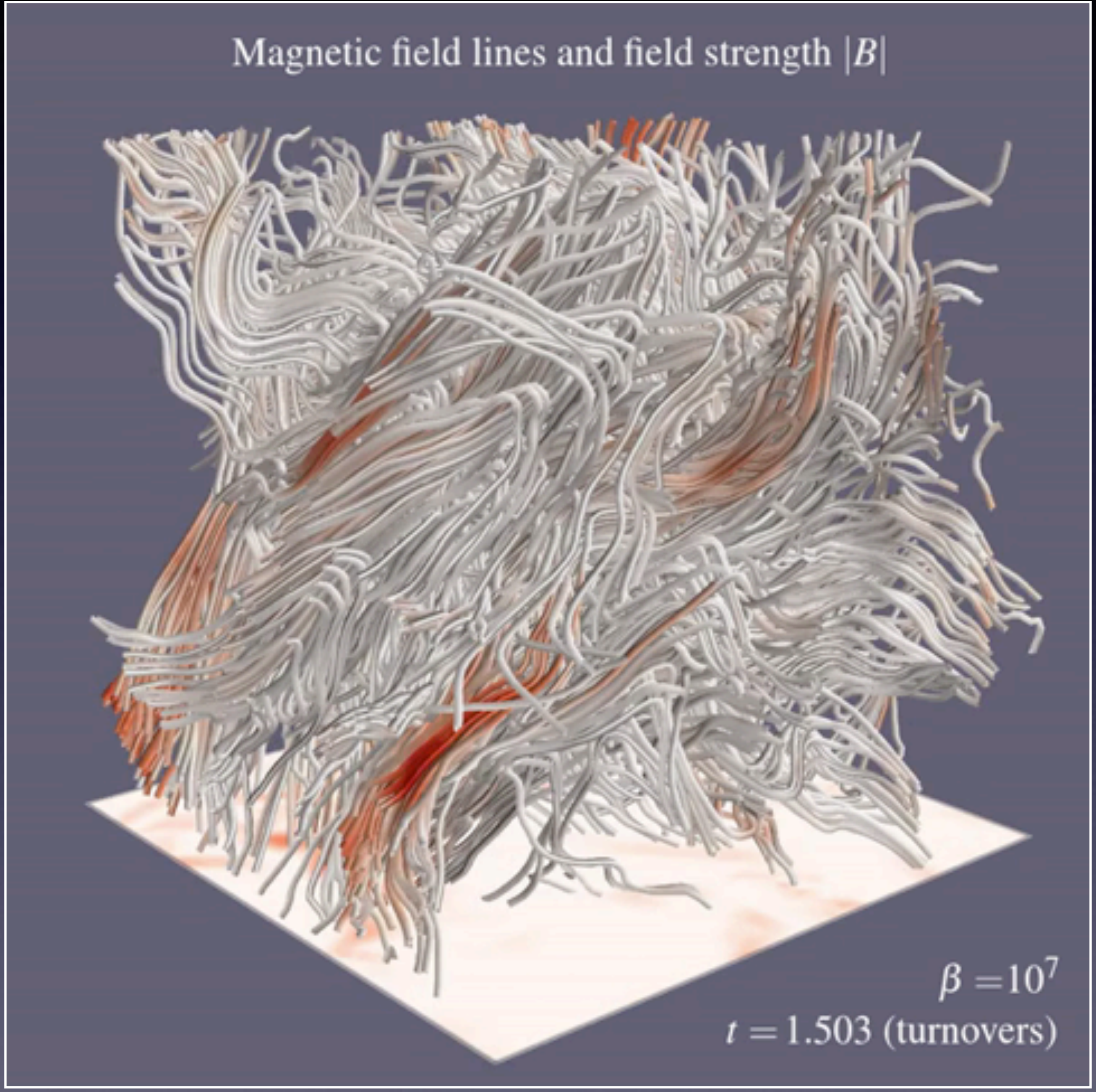
# Exploring the magnetization transition

- Four simulations with same resistivity and input power, but different initial values of $\beta$




- Magnetic growth appears to self-accelerate

# Magnetization transition

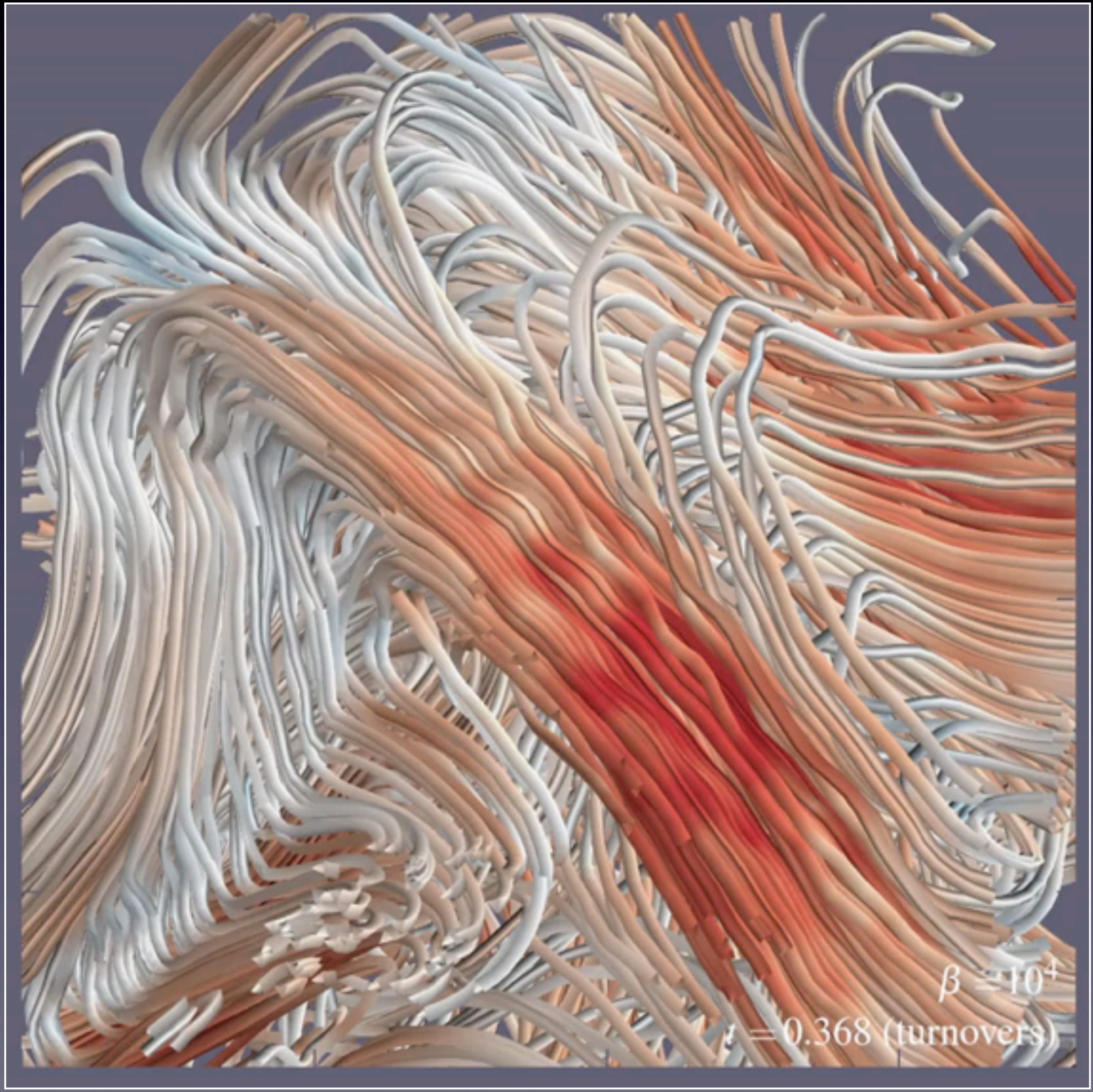



$$\beta = 10^7$$

$$\rho_i/L \simeq 0.5$$

- No scale-separation between stirring and kinetic scales !

# Magnetized regime

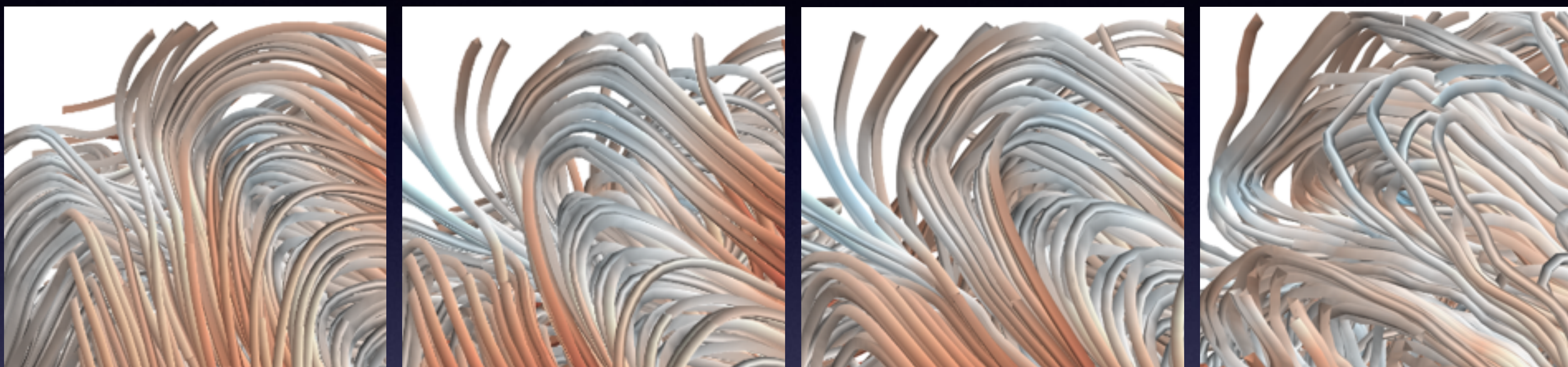



$\beta = 10^4$

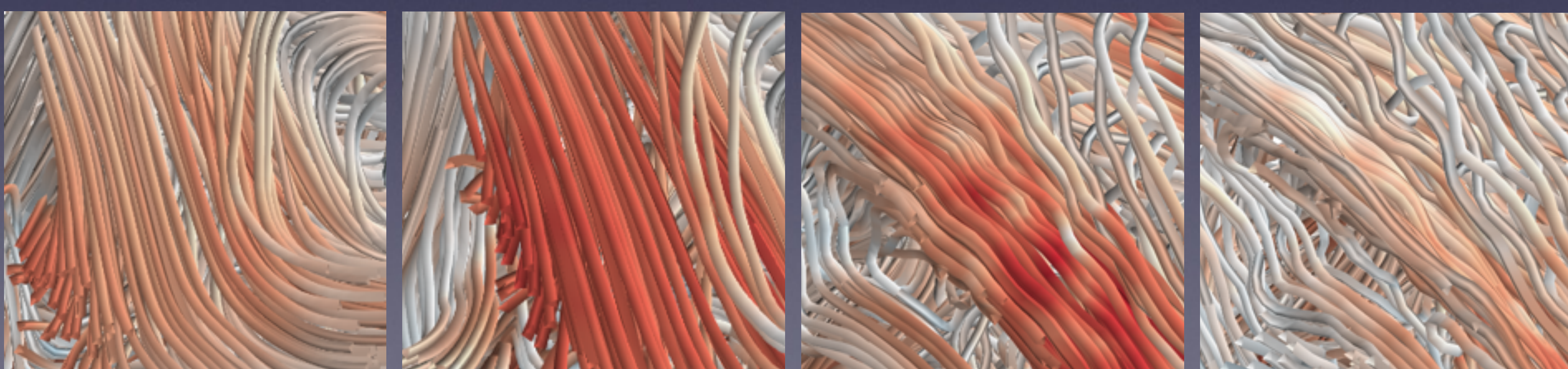
$\rho_i/L \simeq 0.02$

# Magnetized regime

- Firehose instability in strong-field curvature regions

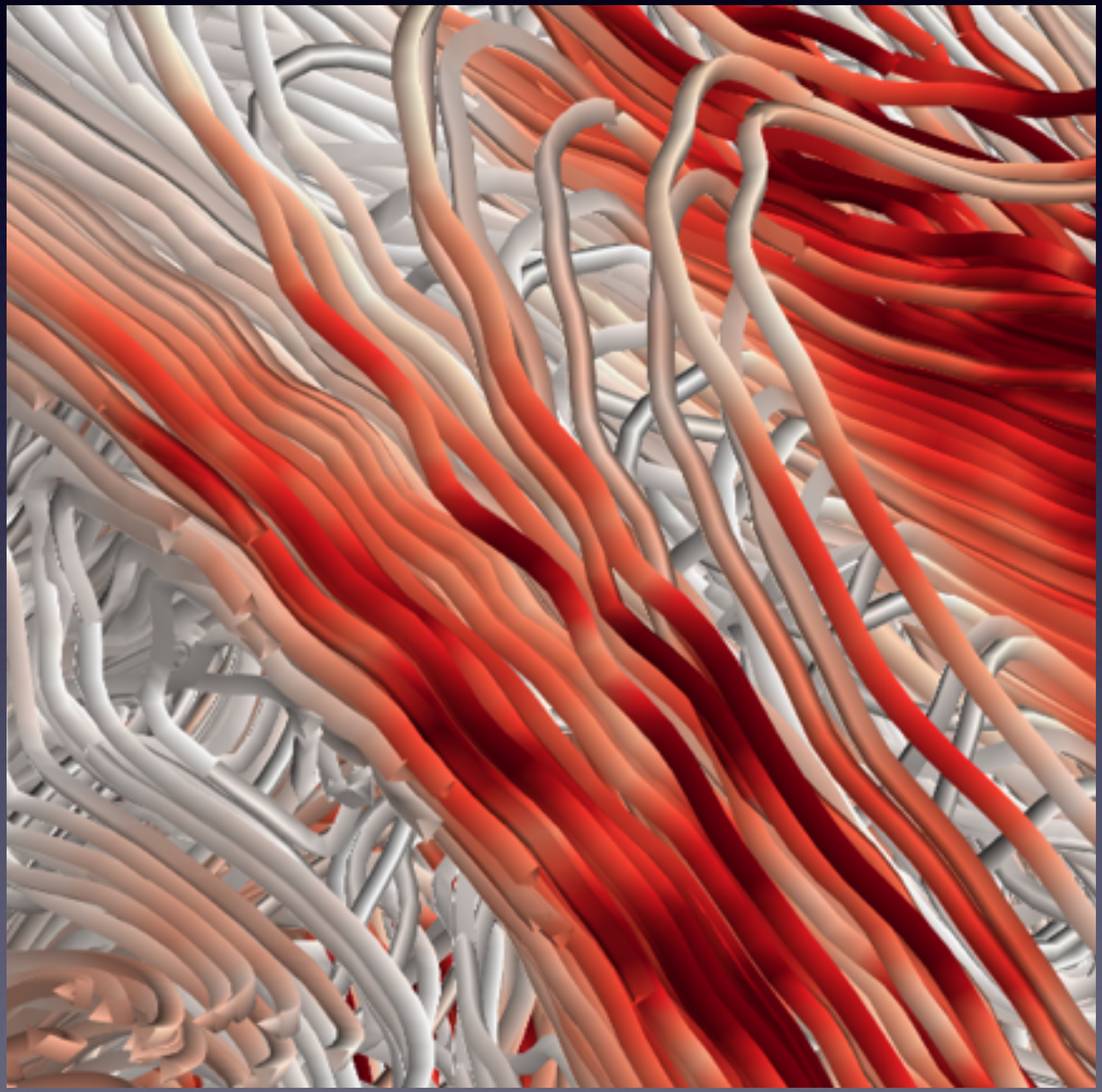


- Bubbly mirror fluctuations in field-stretching regions

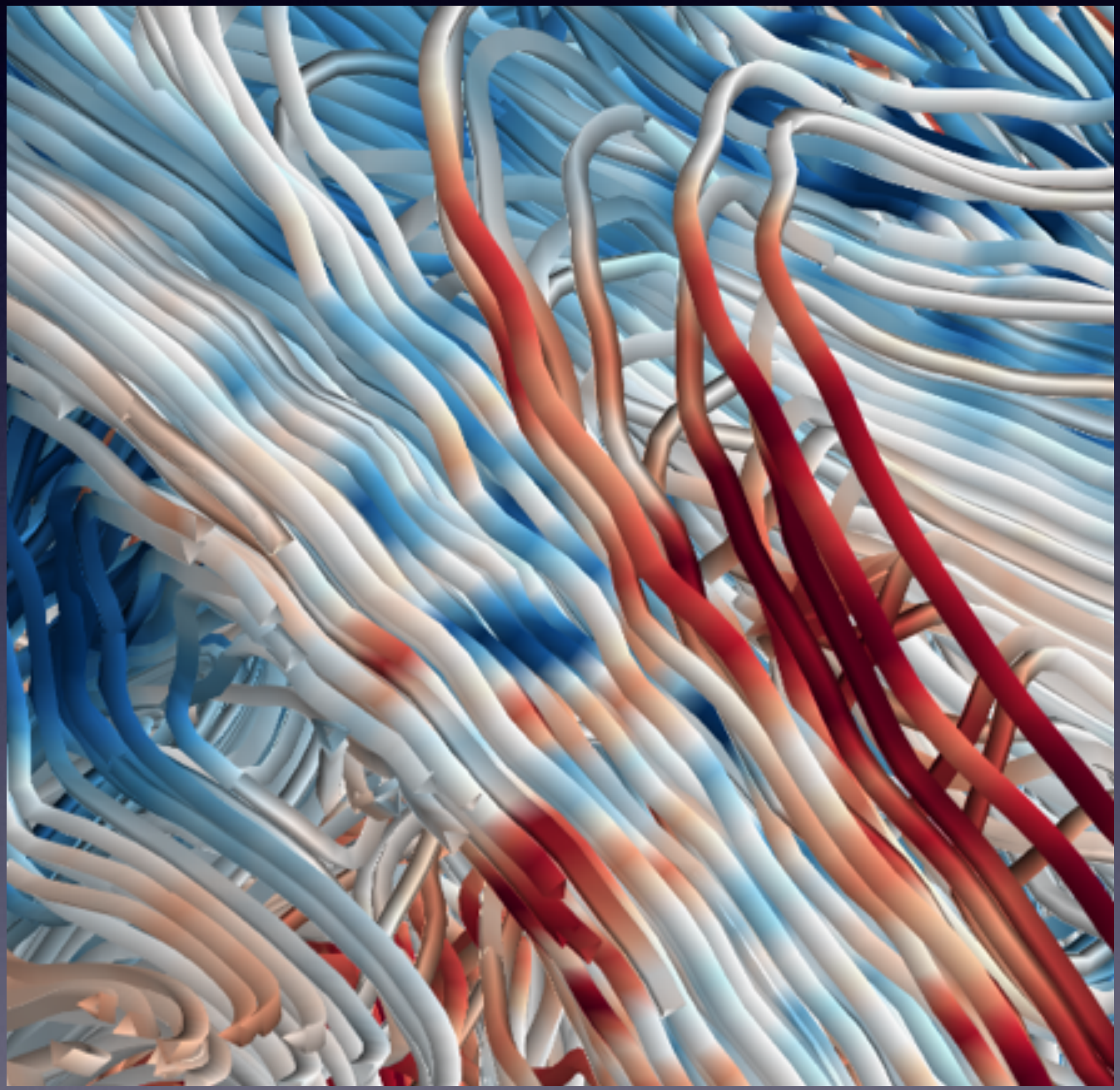


# Magnetized regime

- Mirror structures: magnetic depressions and overdensities



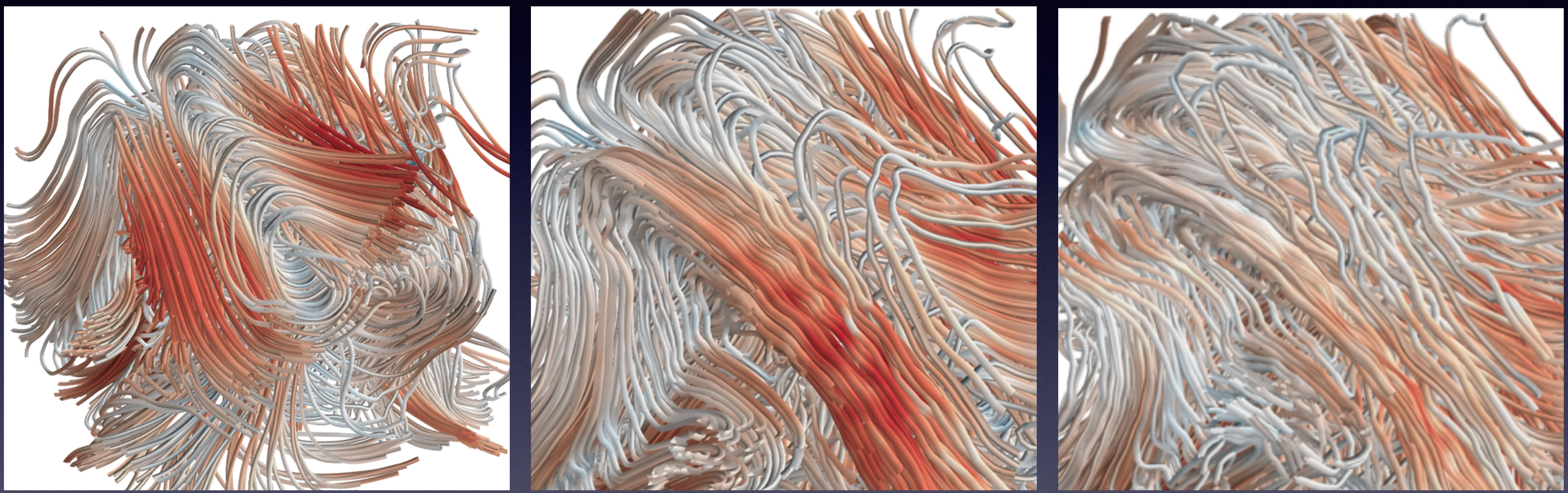
Magnetic strength



Density fluctuations

# Magnetized regime

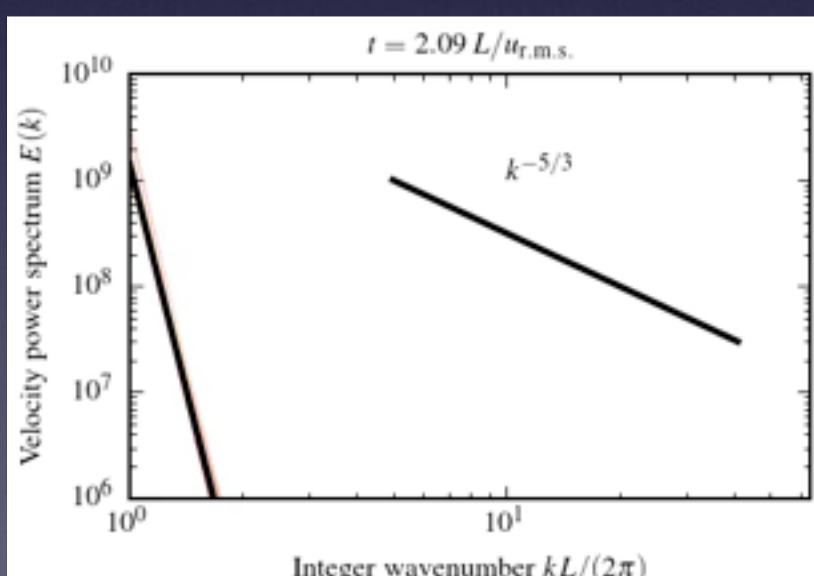
- Pressure anisotropy relaxation



- Current limitations
  - Resolution: cannot go much further at $64^3$ x $51^3$
  - Simulations on longer timescales needed: expensive due to tiny timesteps

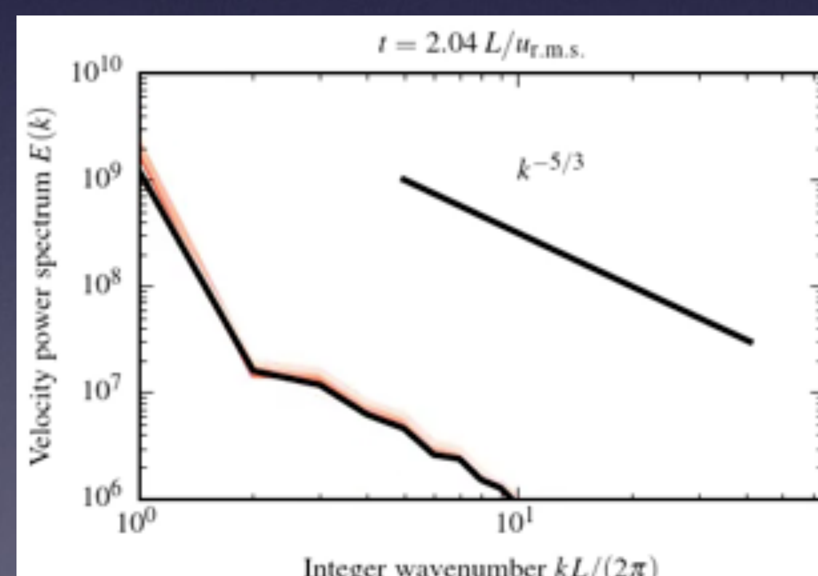
# Ideas on dynamo self-acceleration

- Several "nonlinear" effects possible
  - Dynamo growth entangled with kinetic mode growth
  - Net nonlinear feedback of kinetic modes (see Matt Kunz's talk)
  - Flow viscosity decreases at magnetisation transition, eddies with larger rates of strains are generated



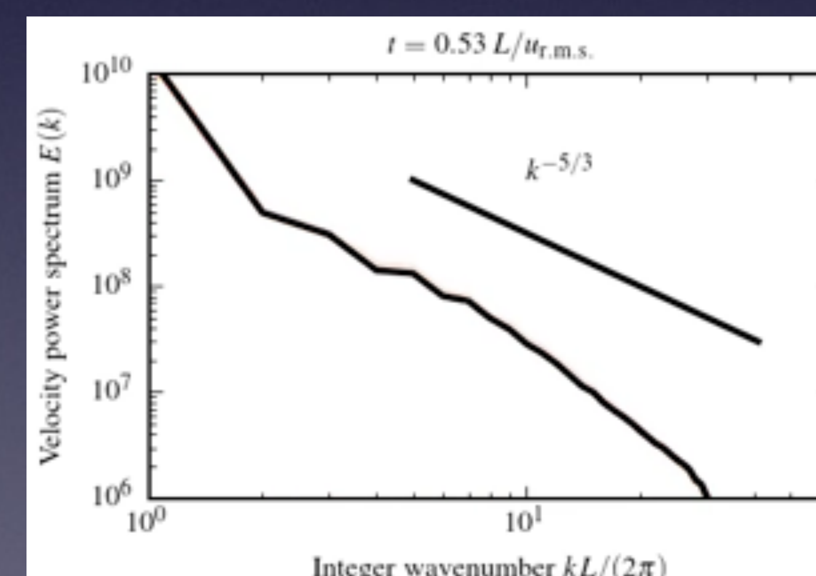

$\beta = 10^{10}$

$\rho_i/L \simeq 16$




$\beta = 10^{7}$

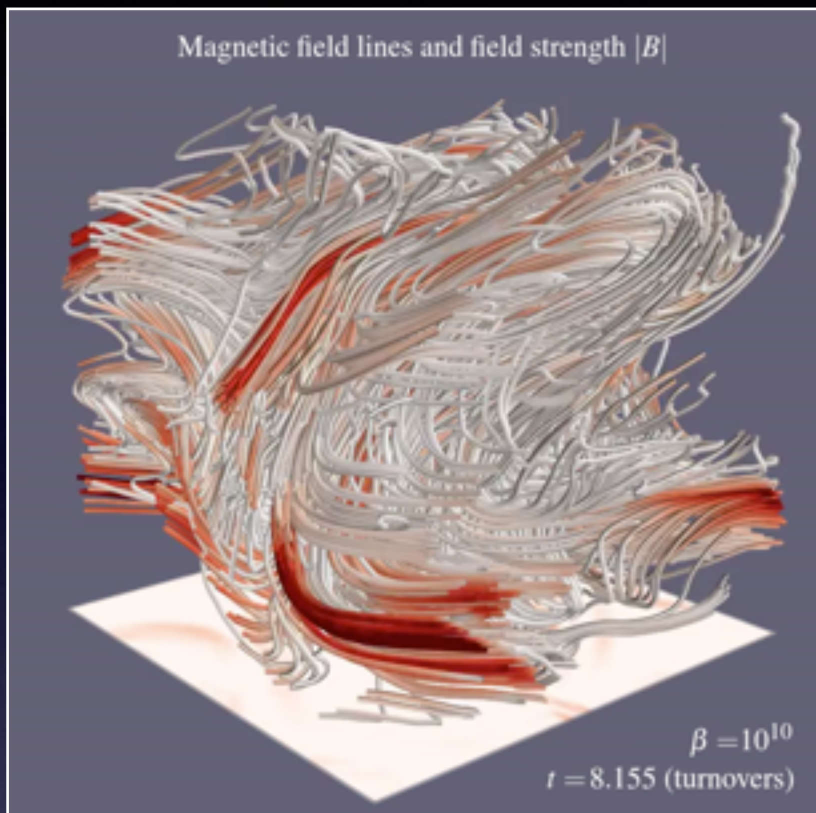
$\rho_i/L \simeq 0.5$

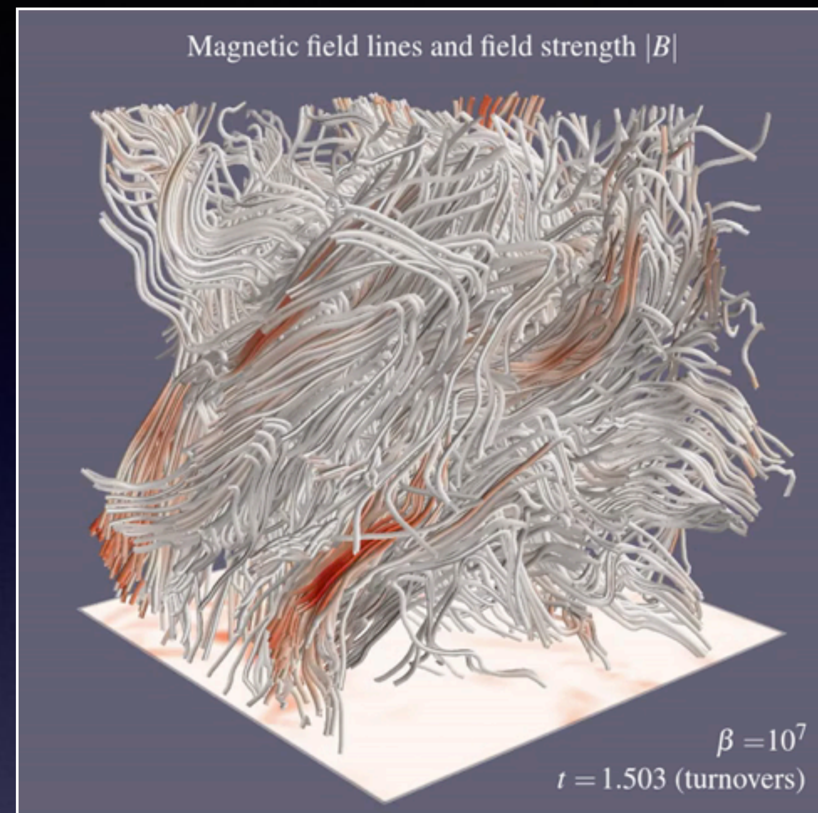


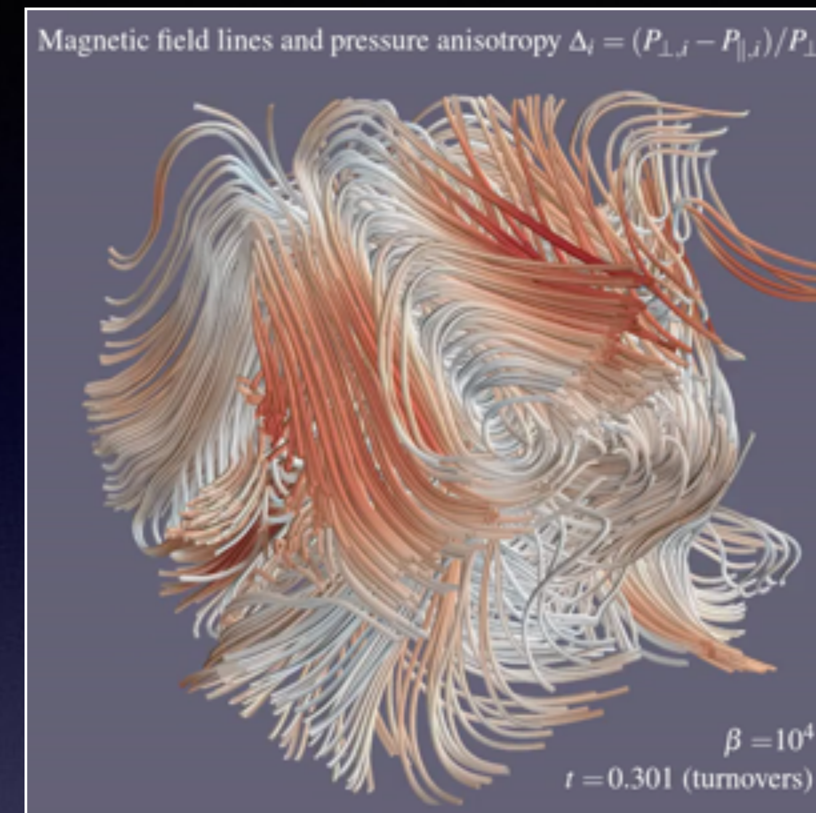

$\beta = 10^{4}$

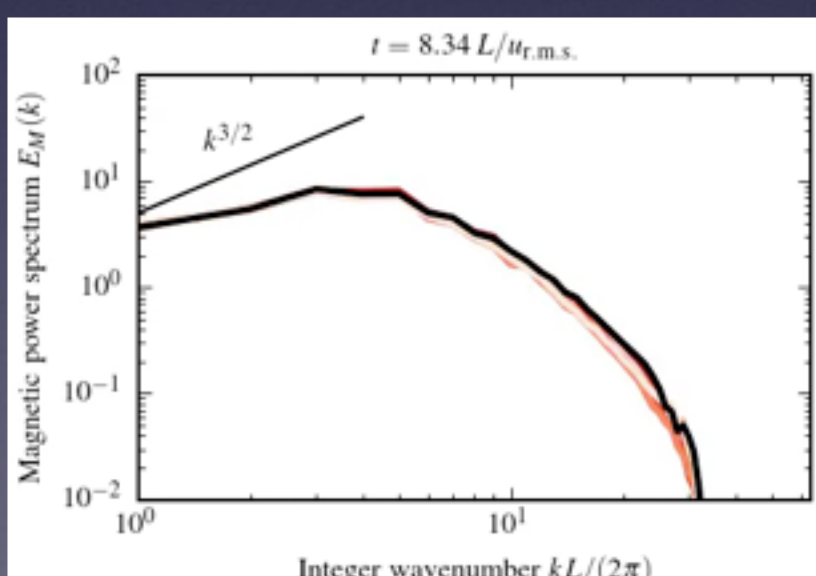
$\rho_i/L \simeq 0.02$

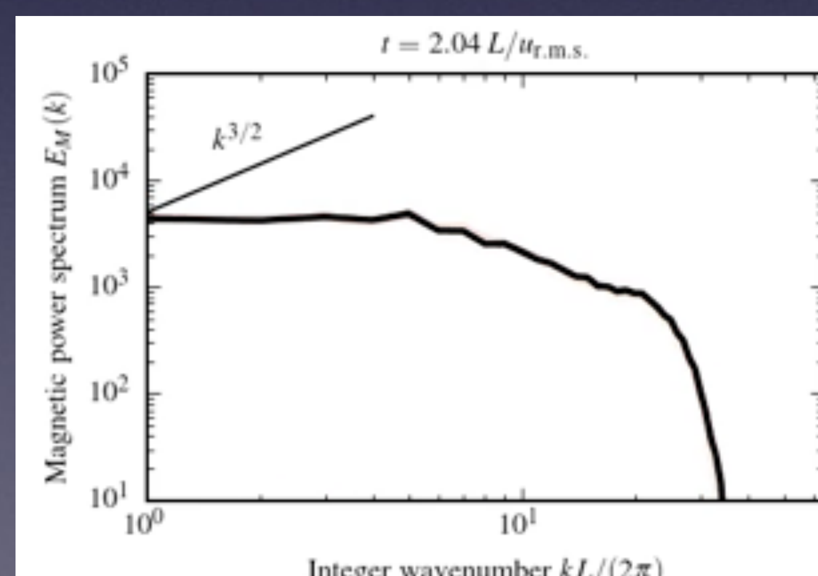
# Magnetic spectra

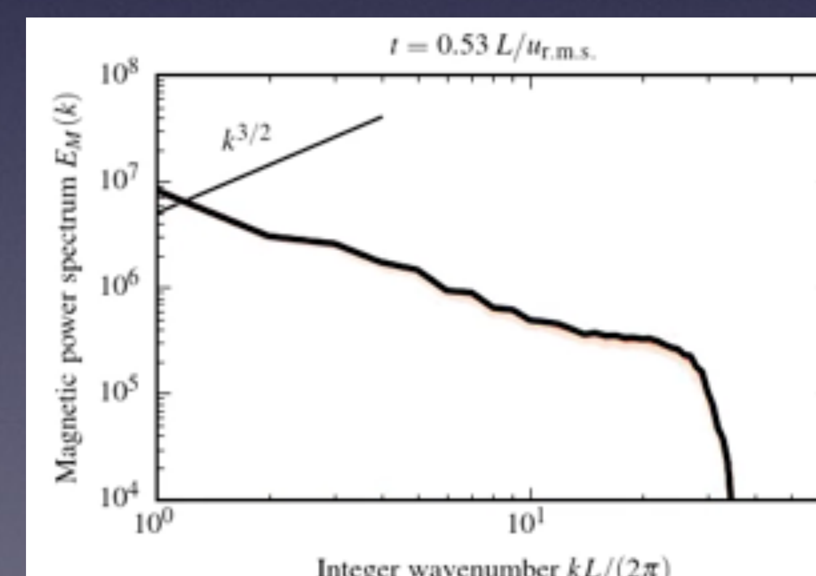



$\beta = 10^{10}$
$\rho_i/L \simeq 16$

$\beta = 10^{7}$
$\rho_i/L \simeq 0.5$

$\beta = 10^{4}$
$\rho_i/L \simeq 0.02$

# Main results and conclusions

- Dynamo in an unmagnetized collisionless plasma is possible
  - Reminiscent of turbulent large Pm MHD dynamo
- Growth self-accelerates as the plasma gets magnetized
- Dynamo and kinetic instabilities become entangled in the magnetized regime
  - Firehose instability in regions of strong field-curvature (negative $\Delta_i$)
  - Mirror instability in regions of field amplification (positive $\Delta_i$)
  - Evolution towards pressure-anisotropy-relaxed state
- Dynamo appears to be a viable mechanism to amplify magnetic field to equipartition in weakly collisional extragalactic plasmas

# Perspectives

- Many interesting questions as yet largely unanswered
  - Magnetization stage really hard to understand: no scale-separation
  - Effective impact of (nonlinear) kinetic instabilities on magnetic growth
  - Dynamical saturation
  - Helical collisionless dynamo
- Future progress rests on
  - Higher-resolution simulations integrated over longer times
  - Experiments
  - Theory !
  - Radio and X-ray astronomy: SKA, Astro-H, Athena+

# (Not) The end